

A Theoretical, Neural-Network Assisted Study of Methane Formation under Astrophysical Conditions

R. Schömig,¹ G. Molpeceres,¹ and J. Kästner¹

¹*Institute for Theoretical Chemistry, University of Stuttgart, Germany*

Dense molecular clouds are cold regions ($\sim 10\text{-}20$ K) in the interstellar medium (ISM) that accommodate dust grains coated with ice, usually amorphous. The chemical components of those ices are H_2O , CO , CO_2 , CH_3OH , etc. with strongly fluctuating compositions. These surfaces play a central role in the chemistry of the ISM and therefore in the evolution of molecular clouds, acting as templates of atoms, radicals and molecules, to meet on top of them [1]. Additionally these ices help to dissipate the energy of the reactions occurring on them. Hence, these surfaces enable the synthesis of a variety of chemical compounds that would not be accessible by gas phase reactions alone. In this work, we studied the formation of CH_4 by investigating the reaction of a CH_3 radical with H on the surface of amorphous solid water (ASW). Among the things we analyzed is the energy redistribution vs chemical desorption [2, 3]. A periodic model of ASW was used to mimic the reaction. As a first step, we have obtained the binding energy (BE) distribution of CH_3 on the surface with values between -500 and -4000 K. The reaction was modeled by placing CH_3 on binding sites with high, medium and low BEs for both axial and planar orientation of the CH_3 radical on the surface. Later, hydrogen atoms have been placed spherically around the CH_3 radical to sample many possible trajectories. Subsequently, molecular dynamics simulations for the reaction were carried out at a temperature of 10 K. The potential driving the dynamics is a neural network potential trained on DFT energies and forces [4, 5]. We have found that the initial orientation of the CH_3 radical hardly plays any role on the reaction. On the contrary, we found huge differences in the evolution of the reaction as a function of the CH_3 binding site and the initial orientation of the H with respect to the methyl radical.

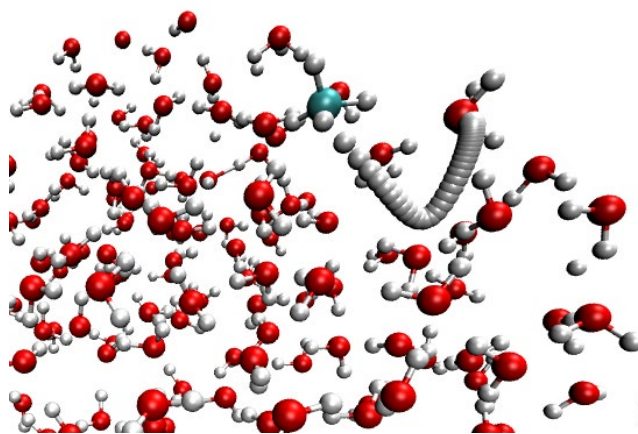


Figure 1: Trajectory of the hydrogen atom approaching the methyl radical to form methane during the MD simulation in increments of 2 fs per hydrogen frame.

References

- [1] Cuppen, H., Walsh, C., Lamberts, T., et al. 2017, *Space Sci. Rev.* **212**, 1
- [2] Garrod, R. T., Wakelam, V., Herbst, E. 2007, *Astron & Astrophys.* **467**, 1103
- [3] Pantaleone, S., Enrique-Romero, J., Ceccarelli, C., et al. 2021, *Astrophys. J.* **917**, 49
- [4] Molpeceres, G., Zaverkin, V., & Kästner, J. 2020, *Mon. Not. R. Astron. Soc.* **499**, 1373
- [5] Zaverkin, V., & Kästner, J. 2020, *J. Chem. Theory Comput.* **16**, 5410